

Forecast on lepton asymmetry from future CMB experiments

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ABSTRACT

We consider a cosmological lepton asymmetry in the form of neutrinos and impose new expected sensitivities on such asymmetry through the degeneracy parameter (ξ_ν) by using some future CMB experiment configurations, such as CORE and CMB-S4. Taking the default scenario with three neutrino states, we find $\xi_\mu = 0.05 \pm 0.10 (\pm 0.04)$, from CORE (CMB-S4) at 95 per cent CL, respectively. Also, within this scenario, we evaluate the neutrino mass scale, obtaining that the normal hierarchy mass scheme is privileged. Our results are an update concerning on the cosmological lepton asymmetry and the neutrino mass scale within this context, from which can bring a perspective on the null hypothesis for ξ_ν (and its effects on ΔN_{eff}), where perhaps, ξ_ν may take a non-null value up to 95 per cent CL from future experiments such as CMB-S4. Sensitivity results for CMB-S4 obtained here not including all expected systematic errors.

Key words: cosmic background radiation – cosmological parameters – cosmology: theory.

1 INTRODUCTION

The lepton asymmetry of the Universe, represented by neutrinos and antineutrinos, is nowadays one of the most weakly constrained cosmological parameter. Although the baryon number asymmetry is well measured from cosmic microwave background (CMB) constraints concerning the baryon density, the lepton asymmetry could be larger by many orders of magnitude and not of the same order as expected by the big bang Nucleosynthesis (BBN) considerations (Barenboim, Kinney & Park 2017). The presence of a large lepton asymmetry can be considered as an excess of the neutrinos on antineutrinos or vice versa, which can be a requirement due to both the charge neutrality of the Universe and it is possibly hidden in the cosmic neutrino background ($C\nu B$) and it can be imprinted on cosmological observation. For instance, from CMB anisotropy (Castorina et al. 2012; Oldengott & Schwarz 2017). The large neutrino asymmetries have consequences in the early Universe phase transitions, cosmological magnetic fields, and dark matter relic density (see Schwarz & Stuke 2009; Semikoz, Sokoloff & Valle 2009; Stuke, Schwarz & Starkman 2012; Barenboim & Park 2017, for more details). Other effects due to the asymmetric leptonic can be considered as changes in the decoupling temperature of $C\nu B$ (Freese, Kolb & Turner 1983; Kang & Steigman 1992), the time equivalence between the energy densities of radiation and matter, the production of

primordial light elements at BBN (Sarkar 1996), an excess in the contribution of the total radiation energy density and the expansion rate of the Universe (Giusarma et al. 2011; Allison et al. 2015), photon decoupling (Lesgourgues & Pastor 1999), among others. These changes can affect the evolution of the matter density perturbations in the Universe, which has effects not only on the CMB anisotropies, but also on the formation, evolution, and distribution of the large-scale structure (LSS) of the Universe (Lesgourgues et al. 2013; Allison et al. 2015). The effects of the cosmological neutrinos on both the CMB and LSS are only gravitational, since they are decoupled (free streaming particles) at the time of recombination and structure formation. The LSS formation is more sensitive to the neutrino masses than CMB. The increasing of the structure is driven by the cosmic expansion and self-gravity of matter perturbations, both affected by the massive neutrinos. Nevertheless, the relic neutrino slows down the growth of structures due to its high thermal speed, leading to a suppression of the total matter power spectrum (Ali-Haïmoud & Bird 2013). On the other hand, the gravitational lensing of CMB and the integrated Sachs–Wolf effect are also modified by the presence of massive neutrinos (Abazajian et al. 2015). The effect of massive neutrinos in the non-linear growth structure regime has been recently studied by Zeng et al. (2018).

The neutrinos properties are very important in the determination of the dynamics of the Universe inferring direct effects on cosmological sources and consequently in the estimation of cosmological parameters (see Dolgov 2002; Lesgourgues & Pastor 2006; Abazajian et al. 2015; Wang et al. 2016; Lorenz, Calabrese &

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Alonso 2017; Vagnozzi et al. 2017, 2018b; Yang et al. 2017; Choudhury & Choubey 2018; Di Valentino et al. 2018a; Li et al. 2018; Wang et al. 2018; Wang, Wang & Xia 2018). The parameters that characterize the neutrinos effects on cosmological probes are the total neutrino mass Σm_ν and the effective number of species N_{eff} . Altogether, the updated constraint on the neutrino mass scale is $\Sigma m_\nu < 0.12$ eV at 95 per cent CL and $N_{\text{eff}} = 2.99 \pm 0.17$ at 95 per cent CL from the final full-mission Planck measurements of the CMB anisotropies (Planck Collaboration VI 2018). In the case of the three active neutrino flavours with zero asymmetries and a standard thermal history, the value of effective number of species is the well-known $N_{\text{eff}} = 3.046$ (Lesgourgues et al. 2013) and an improved calculation $N_{\text{eff}} = 3.045$ (de Salas & Pastor 2016), but the presence of neutrino asymmetries can increase this number without the need to introduce new relativistic species. In general terms, any excess over this value can be parametrized through $\Delta N_{\text{eff}} = N_{\text{eff}} - 3.046$, which in principle is assumed to be some excess of the number of relativistic relics degrees of freedom, known in the literature as dark radiation (see Nunes & Bonilla 2018, for recent constraints on ΔN_{eff}).

Finally, and more important to our work, is to consider the aforementioned cosmological lepton symmetry, which is another natural extension on the neutrino physics properties. This property is usually parametrized by the so-called degeneracy parameter $\xi_\nu = \mu_\nu/T_{\nu 0}$, where μ_ν is the neutrino chemical potential and $T_{\nu 0}$ is the current temperature of the relic neutrino spectrum $T_{\nu 0} \approx 1.9$ K. We can assign to chemical potentials a label of its eigenstates of mass, such that $\{u_i\}$ is for neutrinos and $\{-u_i\}$ for antineutrinos. If the neutrinos are Majorana particles, then they must have $u_i = 0$, and if $u_i \neq 0$, neutrinos are Dirac fermions thus, evidence on the null hypothesis is necessary to help to solve this question (Mangano et al. 2012). The difference between $\{\xi_i\}$ and $\{-\xi_i\}$ determines the asymmetry between the density of neutrinos and antineutrinos. Then, the presence of a relevant and non-zero ξ_ν have some cosmological implications (Hu et al. 1995; Kinney & Riotto 1999; Lesgourgues & Pastor 1999; Abazajian, Beacom & Bell 2002; Wong 2002; Lattanzi, Ruffni & Vereshchagin 2005; Serpico & Raffelt 2005; Ichiki, Yamaguchi & Yokoyama 2007; Popa & Vasile 2008; Simha & Steigman 2008; Schwarz & Stuke 2013; Caramete & Popa 2014; Oldengott & Schwarz 2017). From the particle physics point of view, to measure the lepton asymmetry of the Universe is crucial to understand some of the particle physics processes that might have taken place in early Universe at high energies, including the better constraint on models for the creation of matter–antimatter asymmetry in the Universe (Affleck & Dine 1985; Casas, Cheng & Gelmini 1999; Canetti, Drewes & Shaposhnikov 2012). The tightest constraints on lepton asymmetry at present are commonly based on a combination of CMB data via constraints on the baryon density and measurements of the primordial abundances of light elements (Simha & Steigman 2008; Mangano et al. 2012; Cooke et al. 2014).

In this work, our main target is to obtain new and precise limits on the cosmological lepton asymmetry, measured in terms of the degeneracy parameter ξ_ν , as well as the neutrino mass scale, taking as a basis, the configurations of future CMB experiments such as CMB-CORE and CMB-S4.

This paper is organized as follows. In the next section, we briefly comment on the $C\nu B$ and the cosmological lepton asymmetry. In Section 3, we present the methodology used to obtain the forecasts from CMB-CORE and CMB-S4 experiments. In Section 4, we present our results and discussions. Finally, in Section 5 our final considerations.

2 COSMIC NEUTRINO BACKGROUND AND NEUTRINO ASYMMETRY

The current contribution of neutrinos to the energy density of the Universe is given by,

$$\rho_\nu = 10^4 h^2 \Omega_\nu \text{ eV cm}^{-3}, \quad (1)$$

where $\Omega_\nu = \rho_\nu/\rho_{\text{crit}}$ is the neutrino energy density in units of critical density. As usual, relativistic neutrinos contribute to the total energy density of radiation ρ_r , typically parametrized as

$$\rho_r = (\rho_\gamma + \rho_\nu) = \left(1 + \frac{7}{8} \left(\frac{4}{11}\right)^{4/3} N_{\text{eff}}\right) \rho_\gamma, \quad (2)$$

where ρ_γ is the energy density of photons, the factor 7/8 is due to the neutrinos that are fermions and $N_{\text{eff}} = 3.046$ the value of the effective number of neutrinos species in the standard case, with zero asymmetries and no extra relativistic degrees of freedom. Neutrinos become non-relativistic when their average momentum falls below their mass. At the very early Universe, neutrinos and antineutrinos of each flavour ν_i ($i = e, \mu, \tau$) behave like relativistic particles. Both the energy density and pressure of one species of massive degenerate neutrinos and antineutrinos are described by (let us use here the unit system where $\hbar = c = k_B = 1$)

$$\rho_{\nu_i} + \rho_{\bar{\nu}_i} = T_{\nu_i}^4 \int \frac{d^3 q}{2(\pi)^3} E_{\nu_i}(f_{\nu_i}(q) + f_{\bar{\nu}_i}(q)) \quad (3)$$

and

$$3(p_{\nu_i} + p_{\bar{\nu}_i}) = T_{\nu_i}^4 \int \frac{d^3 q}{2(\pi)^3} \frac{q^2}{E_{\nu_i}} (f_{\nu_i}(q) + f_{\bar{\nu}_i}(q)), \quad (4)$$

where $E_{\nu_i}^2 = q^2 + a^2 m_{\nu_i}$ is one flavour neutrino/antineutrino energy and $q = ap$ is the comoving momentum. The functions $f_{\nu_i}, f_{\bar{\nu}_i}$ are the Fermi–Dirac phase-space distributions given by

$$f_{\nu_i}(q) = \frac{1}{e^{E_{\nu_i}/T_{\nu_i} - \xi_\nu} + 1}, f_{\bar{\nu}_i}(q) = \frac{1}{e^{E_{\bar{\nu}_i}/T_{\nu_i} + \xi_\nu} + 1}, \quad (5)$$

where $\xi_\nu = \mu_\nu/T_{\nu 0}$ is the neutrino degeneracy parameter and μ_ν is the neutrino chemical potential. At the early Universe, we assumed that neutrinos–antineutrinos are produced in thermal and chemical equilibrium. Their equilibrium distribution functions have been frozen from the time of decoupling to the present. Then, as the chemical potential μ_ν scales as T_ν , the degeneracy parameter ξ_ν remains constant and it is different from zero if a neutrino–antineutrino asymmetry has been produced before the decoupled. The energy of neutrinos changes according to cosmological redshift after decoupling, which is a moment when they are still relativistic. The neutrino degeneracy parameter is conserved and the presence of a significant and non-null ξ_ν have some cosmological implication through the evolution of the universe, such as BBN, photon decoupling, and LSS, among others (see Hu et al. 1995; Kinney & Riotto 1999; Lesgourgues & Pastor 1999; Abazajian et al. 2002; Wong 2002; Lattanzi et al. 2005; Serpico & Raffelt 2005; Ichiki et al. 2007; Popa & Vasile 2008; Simha & Steigman 2008; Caramete & Popa 2014). If ξ_ν remains constant, finite, and non-zero after decoupling, then it could lead to an asymmetry on the neutrinos and antineutrinos given by

$$\eta_\nu \equiv \frac{n_{\nu_i} - n_{\bar{\nu}_i}}{n_\gamma} = \frac{1}{12\zeta(3)} \sum_i y_{\nu 0} (\pi^2 \xi_{\nu_i} + \xi_{\nu_i}^3), \quad (6)$$

where $n_{\nu_i} (n_{\bar{\nu}_i})$ is the neutrino (antineutrino) number density, n_γ is the photon number density, $\zeta(3) \approx 1.20206$, and $y_{\nu 0}^{1/3} = T_{\nu 0}/T_{\gamma 0}$ is the ratio of neutrino and photons temperature to the present, where $T_{\gamma 0}$ is the temperature of the CMB ($T_{\gamma 0} = 2.726$ K).

As we have mentioned above, the neutrino asymmetry can produce changes in the expansion rate of the Universe at early times, which can be expressed as an excess in N_{eff} in the form

$$\Delta N_{\text{eff}} = \frac{15}{7} \sum_i \left[2 \left(\frac{\xi_{\nu_i}}{\pi} \right)^2 + \left(\frac{\xi_{\nu_i}}{\pi} \right)^4 \right]. \quad (7)$$

In what follows, let us impose expected sensitivities on ξ by taking predictions from some future CMB experiments.

3 METHODOLOGY

Let us predict the ability of future CMB experiments to constrain the neutrino lepton asymmetry as well as the neutrino mass scale. We follow the common approach already used (see e.g. Di Valentino et al. 2018b; Finelli et al. 2018), on mock data for some possible future experimental configurations, assuming a fiducial flat Λ CDM model compatible with the Planck 2018 results. We have used the publicly available Boltzmann code CLASS (Blas, Lesgourgues & Tram 2011) to compute the theoretical CMB angular power spectra C_l^{TT} , C_l^{TE} , C_l^{EE} for temperature, cross temperature-polarization, and polarization. Together with the primary anisotropy signal, we have also taken into account informations from CMB weak lensing, considering the power spectrum of the CMB lensing potential C_l^{PP} . The missions BB are clearly sensitive also to the BB lensing polarization signal, but here we take the conservative approach to not include it in the forecasts.

In our simulations, we have used an instrumental noise given by the usual expression

$$N_l = w^{-1} \exp(l(l+1)\theta^2/8 \ln(2)), \quad (8)$$

where θ is the experimental full width at half-maximum angular resolution, w^{-1} is the experimental power noise expressed in $\mu\text{K-arcmin}$. The total variance of the multipoles a_{lm} is therefore given by the sum of the fiducial C_l 's with the instrumental noise N_l .

The simulated experimental data are then compared with a theoretical model assuming a Gaussian likelihood \mathcal{L} given by

$$-2 \ln \mathcal{L} = \sum_l (2l+1) f_{\text{sky}} \left(\frac{D}{|\bar{C}_l|} + \ln \frac{|\hat{C}_l|}{|\bar{C}_l|} - 3 \right), \quad (9)$$

where \bar{C}_l and \hat{C}_l are the assumed fiducial and theoretical spectra plus noise, and $|\bar{C}_l|$ and $|\hat{C}_l|$ are the determinants of the theoretical and observed data covariance matrices given by

$$|\bar{C}_l| = \bar{C}_l^{TT} \bar{C}_l^{EE} \bar{C}_l^{PP} - (\bar{C}_l^{TE})^2 \bar{C}_l^{PP} - (\bar{C}_l^{TP})^2 \bar{C}_l^{EE}, \quad (10)$$

$$|\hat{C}_l| = \hat{C}_l^{TT} \hat{C}_l^{EE} \hat{C}_l^{PP} - (\hat{C}_l^{TE})^2 \hat{C}_l^{PP} - (\hat{C}_l^{TP})^2 \hat{C}_l^{EE}, \quad (11)$$

D is defined as

$$D = \hat{C}_l^{TT} \bar{C}_l^{EE} \bar{C}_l^{PP} + \bar{C}_l^{TT} \hat{C}_l^{EE} \bar{C}_l^{PP} + \bar{C}_l^{TT} \bar{C}_l^{EE} \hat{C}_l^{PP} - \bar{C}_l^{TE} (\bar{C}_l^{TE} \hat{C}_l^{PP} + 2 \hat{C}_l^{TE} \bar{C}_l^{PP}) - \bar{C}_l^{TP} (\bar{C}_l^{TP} \hat{C}_l^{EE} + 2 \hat{C}_l^{TP} \bar{C}_l^{EE}), \quad (12)$$

and finally f_{sky} is the sky fraction sampled by the experiment after foregrounds removal.

In Table 1 we have summarized the experimental specifications for CMB-CORE and CMB-S4 data. Forecast based on future CMB experiments to probe neutrinos properties were also investigated in Capparelli et al. (2018), Brinckmann et al. (2018), and Mishra-Sharma, Alonso & Dunkley (2018). Specifically for CMB-S4, we use an experimental specification different from that presented in

Mishra-Sharma et al. (2018) and Abazajian et al. (2016) ($l_{\text{max}} = 5000$), where we avoid systematic errors at highest angular resolution ($l > 3000$) produced by noise due to information from CMB weak lensing, which could produce more optimistic results than in previous analyses.

4 RESULTS

We have used the public and available CLASS (Blas et al. 2011) and Monte Python (Audren et al. 2013) codes concerning the model considered in this work, where we introduced the ξ_ν corrections on N_{eff} defined in equation (7) in CLASS code. We have considered one massive and two massless neutrino states, as standard in the literature, and we fixed the mass ordering to the normal hierarchy with the minimum masses $\sum m_\nu = 0.06$ eV and the expected sensitivities obtained on the total neutrino mass are essentially independent of neutrino mixing parameters how is it concluded in Castorina et al. (2012).

The individual values of neutrino flavour asymmetries in principle can be different if we take into account the effect of the oscillations and collisions around the epoch of neutrino decoupling, which means that the equations (6) and (7) are not necessarily valid (Pastor, Pinto & Raffelt 2009; Castorina et al. 2012). However, following Oldengott & Schwarz (2017) we assume a single value of ξ_ν , which means that for the values of the neutrino mixing parameters preferred by global fits of oscillations, and in particular that of $\sin^2 \theta_{13}$, the impact of a lepton asymmetry can be approximated by choosing a common value ξ_ν for degeneracy parameters (Dolgov et al. 2002; Wong 2002; Mangano et al. 2012).

On the other hand, in Castorina et al. (2012) it is shown how the addition of flavour oscillations produces strong constraints on the total neutrino asymmetry, whose boundaries are dominated mainly by the limits imposed by BBN tests. However, although the impact that produces the combined effect of BBN and flavour oscillations on these limits is evident, we are more interested in showing the impact of the improvement in the sensitivity of future CMB experiments.

In our forecasts, we have assumed the set of the cosmological parameters:

$$\left\{ 100\omega_b, \omega_{\text{cdm}}, \ln 10^{10} A_s, n_s, \tau_{\text{reio}}, H_0, \sum m_\nu, \xi_\nu \right\}.$$

where the parameters are: baryon density, cold dark matter density, amplitude and slope of the primordial spectrum of metric fluctuations, optical depth to reionization, Hubble constant, neutrino mass scale, and the degeneracy parameter characterizing the degree of leptonic asymmetry, respectively. In the forecast, we assume fiducial values of $\{2.22, 0.119, 3.07, 0.962, 0.05, 68.0, 0.06, 0.05^1\}$, which are assumed from our analysis performed for Planck 2018.

Table 2 shows the constraints on the model baseline imposed by the CORE and S4 experiments. Figs 1 and 2 shows the parametric space for some parameters of interest in our work, from Planck/CORE and CORE/S4 constraints, respectively. From Planck data, we can note that the degeneracy parameter is constrained to $\xi_\nu = 0.05 \pm 0.20$ (± 0.33) at 68 per cent CL and 95 per cent CL, which is a result compatible with the null hypothesis even to 1σ CL. In Oldengott & Schwarz (2017), the authors obtain $\xi_\nu = -0.002_{-0.11}^{+0.114}$ at 95 per cent CL from Planck data. Evidence for

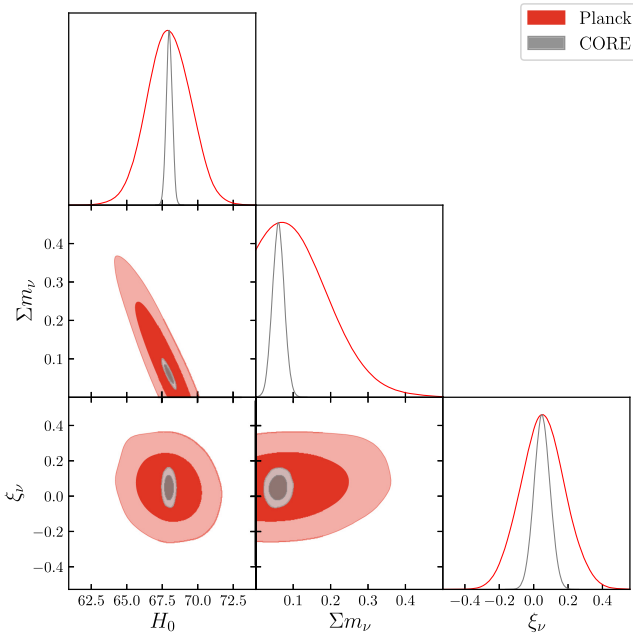
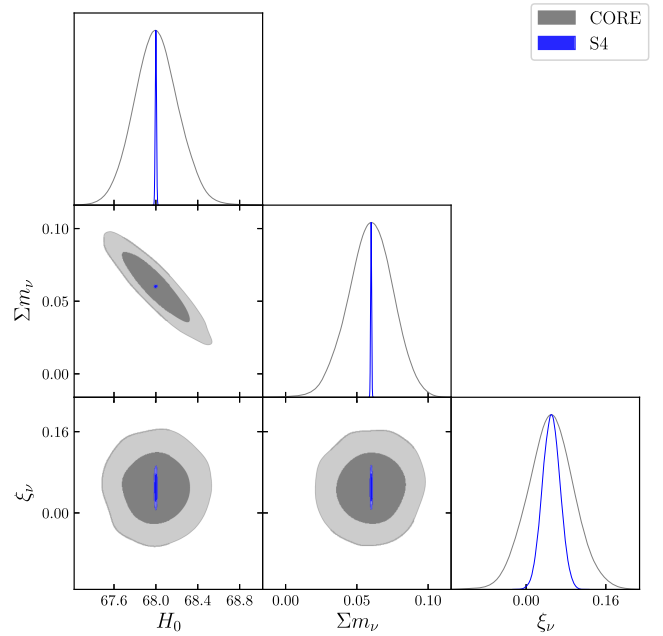
¹This value is in accordance with the results obtained in (Nunes & Bonilla 2018) (Table III) and Castorina et al. (2012).

Table 1. Experimental specifications for CORE and S4 with beamwidth, power noise sensitivities of the temperature, and polarization.

Experiment	Beam	Power noise ($\mu\text{K-arcmin}$)	l_{\min}	l_{\max}	f_{sky}
CMB-CORE	6.0	2.5	2	3000	0.7
CMB-S4	3.0	1.0	50	3000	0.4

Table 2. Summary of the observational constraints from both CORE and S4 experiments. The notation σ (CORE) and σ (S4), represents the 68 per cent CL estimation on the fiducial values from CORE and S4, respectively. The parameter H_0 is in $\text{km s}^{-1} \text{Mpc}^{-1}$ units and $\sum m_\nu$ is in eV units.

Parameter	Fiducial value	σ (CORE)	σ (S4)
$10^2 \omega_b$	2.22	0.000057	0.00012
ω_{cdm}	0.11919	0.00037	0.000093
H_0	68.0	0.32	0.0088
$\ln 10^{10} A_s$	3.0753	0.0056	0.0035
n_s	0.96229	0.0022	0.0054
τ_{reio}	0.055	0.0028	0.00025
$\sum m_\nu$	0.06	0.024	0.00053
ξ_ν	0.05	0.071	0.027


Figure 1. One-dimensional marginalized distribution and 68 per cent CL and 95 per cent CL regions for some selected parameters taking into account Planck and CORE experiments.

Figure 2. One-dimensional marginalized distribution and 68 per cent CL and 95 per cent CL regions for some selected parameters taking into account CORE and S4 experiments.

cosmological lepton asymmetry from CMB data have been found by Caramete & Popa (2014).

On the other hand, the constraints on the degeneracy parameter are close to the null value also within the accuracy achieved by CORE data, $\xi_\nu = 0.05 \pm 0.071$ (± 0.11) at 68 per cent CL and 95 per cent CL, being compatible with the null hypothesis even to 1σ CL, as in the case of Planck data, used in this work. However, with respect to the accuracy obtained by CMB-S4, we find $\xi_\nu = 0.05 \pm 0.027$ (± 0.043) at 68 per cent CL (95 per cent CL), respectively. These constraints can rule out the null hypothesis

up to 2σ CL on ξ_ν . In principle this last result can open the door to the possibility to unveil the physical nature of neutrinos, that is, the neutrinos can be Dirac particles against the null hypothesis and no Majorana particles as established by such hypothesis. However, these results must be firmly established from the point of view of particle physics for example from ground-based experiments such as PandaX-III (Particle And Astrophysical Xenon Experiment III), which to explore the nature of neutrinos, including physical properties such as the absolute scale of the neutrino masses and the aforementioned violation of leptonic number conservation

through Neutrinoless Double Beta Decay (NLDBD)² and whose observation will be a clear signal that the neutrinos are their own antiparticles (for more details see Chen et al. 2017). These results could be available within the first 5–10 yr of the next decade.

In Fig. 1 we can note that there is a high anticorrelation between the neutrinos' masses and H_0 , that will increase the tension between the local and global measures of H_0 , if the masses of the neutrinos increase (therefore they will decrease the value of H_0), such that, constraints on those parameters must be cautiously interpreted until such tension can be better understood. Within the standard base- Λ CDM cosmology, the Planck Collaboration (Planck Collaboration VI 2018) report $H_0 = 67.36 \pm 0.54 \text{ km s}^{-1} \text{ Mpc}^{-1}$ which is about 99 per cent away from the locally measured value $H_0 = 72.24 \pm 1.74 \text{ km s}^{-1} \text{ Mpc}^{-1}$ reported in Riess et al. (2016). We obtain, $H_0 = 68.00 \pm 2.32 (\pm 3.78) \text{ km s}^{-1} \text{ Mpc}^{-1}$ at 68 per cent CL and 95 per cent CL, for our model with Planck data, which at least 2σ can reduce the tension between the global and local value of H_0 . The difference with our results regarding Planck 2018 is due to our extended parameter space. On the other hand, from the Planck data analysis we can note that the neutrino mass scale is constrained to $\sum m_\nu < 0.36 \text{ eV}$ at 95 per cent CL, which is in good agreement with the one obtained by Planck Collaboration, i.e. $\sum m_\nu < 0.24 \text{ eV}$ (Planck Collaboration VI 2018). From the $\sum m_\nu - H_0$ plane, we note that no relevant changes are obtained with respect to the mass splitting, which requires that $\sum m_\nu < 0.1 \text{ eV}$ to rule out the inverted mass hierarchy ($m_2 \gtrsim m_1 \gg m_3$). However, these results start to favour the scheme h of normal hierarchy ($m_1 \ll m_2 < m_3$) which will be evidenced with our results take CORE and S4 predictions. The results from CORE and S4 present considerable improvements with respect to Planck data, see Figs 1 and 2. With respect to neutrino mass scale bounds imposed from CORE and S4 data, we find the limit $0.021 < \sum m_\nu \lesssim 0.1 \text{ eV}$ and $0.05913 < \sum m_\nu \lesssim 0.061 \text{ eV}$ at 95 per cent CL, for CORE and S4, respectively. Thus, unfavourable to inverted hierarchy scheme mass at least at 95 per cent CL in both cases.

In the standard scenario of three active neutrinos and if we consider effects of non-instantaneous decoupling, we have $N_{\text{eff}} = 3.046$, where it is important to make clear that in all the analysis we considered this value as a fixed one. It is well known that the impact of the leptonic asymmetry increase the radiation energy density with the form, $N_{\text{eff}} = 3.046 + \Delta N_{\text{eff}}^{\xi_\nu}$, where $\Delta N_{\text{eff}}^{\xi_\nu}$ is due to the leptonic asymmetry induced via equation (7):

$$\Delta N_{\text{eff}}^{\xi_\nu} = \frac{60}{7} \left(\frac{\xi_\nu}{\pi} \right)^2 + \frac{30}{7} \left(\frac{\xi_\nu}{\pi} \right)^4, \quad (13)$$

where the sum is $i = 1, 2$ only (two massless neutrino states). Without losing of generality, we can evaluate the contribution $\Delta N_{\text{eff}}^{\xi_\nu}$ via the standard error propagation theory. We note that, $\Delta N_{\text{eff}}^{\xi_\nu} = 0.002 \pm 0.019 (\pm 0.030)$ for Planck data, $\Delta N_{\text{eff}}^{\xi_\nu} = 0.0022 \pm 0.0083 (\pm 0.013)$ for CORE data, and $\Delta N_{\text{eff}}^{\xi_\nu} = 0.0022 \pm 0.0045 (\pm 0.0059)$ for S4 data, all limits at 68 per cent and 95 per cent CL. Therefore,

²In nuclear physics Double Beta Decay (DBD) it's a type of radioactive decay process of second-order weak interactions observed experimentally in several isotopes, in which two electrons (positrons) and two antineutrinos (neutrinos) are emitted simultaneously from decaying nucleus (protons into neutrons or vice versa). If neutrinos are Majorana particles, a second DBD mode is possible, where a nucleus can decay again by emitting just two electrons (positrons) without antineutrinos (neutrinos), which are exchanged in the decay of nucleons.

in general lines, we can assert that the contribution from ξ_ν on N_{eff} are very small. But in the case of CMB-S4, even this contribution being very small, it can be non-null.

5 CONCLUSIONS

In this work, we have derived new constraints relative to the lepton asymmetry through the degeneracy parameter by using the CMB angular power spectrum from the Planck data and future CMB experiments like CORE and CMB-S4. We have analysed the impact of a lepton asymmetry on N_{eff} where, as expected, we noticed the existence of very small corrections on ΔN_{eff} , but corrections that cannot be negligible at the level of CMB-S4 experiments, although should be taken into account that sensitivity results obtained from CMB-S4 not including all expected systematic errors, as was previously mentioned (for similar considerations and more detailed information see Allison et al. 2015). Within this cosmological scenario, we have also investigated the neutrino mass scale in combination with the cosmological lepton asymmetry. We have found strong limits on $\sum m_\nu$, where the mass scale for both, CORE and CMB-S4 configurations, are well bound to be $\sum m_\nu < 0.1 \text{ eV}$ at 95 per cent CL, therefore, favouring a normal hierarchy scheme within the perspective adopted here.

As future perspective, it can be interesting to consider a neutrino asymmetry interaction with the dark sector of the Universe, and to see how this coupling can affect the neutrino and dark matter/dark energy properties, as well as to bring possible new corrections on ΔN_{eff} due to such interaction, including properly the effect of flavour oscillations and galaxy bias due to neutrinos, which have been recently targeted in literature. For more details about these last topics and a deeper discussion we cordially invite the reader to Pastor et al. (2009), Castorina et al. (2012), where can be perceived as a proper implementation of the oscillations can significantly improve the constraints on the main physical properties of massive neutrinos and about the neutrino bias galaxy such as in Vagnozzi et al. (2018a); Giusarma et al. (2018), where it is suggested that the proper modelling of the bias parameter is necessary in order to reduce the impact of non-linearities and minimization of systematics, in addition to the need to correct for the neutrino-induced scale-dependent bias, whose correct implementation is still under construction within the community of cosmologists.

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